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Inner horizon of the quantum Reissner-Nordström black holes

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ABSTRACT: We study the nature of the inner Cauchy horizon of a Reissner-Nordström black hole in a quantum context by means of the horizon wave-function obtained from modelling the electrically charged source as a Gaussian wave-function. Our main finding is that there are significant ranges for the black hole mass (around the Planck scale) and specific charge for which the probability of realising the inner horizon is negligible. This result suggests that any semiclassical instability one expects near the inner horizon may not occur in quantum black holes.

KEYWORDS: Models of Quantum Gravity, Black Holes

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1 Introduction

Current literature contains many attempts at quantizing black hole (BH) metrics, which focus on the purely gravitational degrees of freedom, and yield a description of the horizon unrelated to the matter state that sources the geometry [1–4]. However, this point of view might miss important features emerging from the highly non-linear nature of the gravitational interaction. One can think in analogy to the hydrogen atom, and note that its energy levels, which have no classical counterpart, cannot be simply obtained by quantising the free electromagnetic field. A different perspective is taken in the Horizon Wave Function (HWF) formalism [5–7], which is instead based on the quantum version of the Einstein equation relating the size of the gravitational radius (which can be a horizon) to the (quantum) state of matter. This formalism has been applied to several case studies [8, 9], yielding sensible results in agreement with (semi)classical expectations, and there is therefore hope that it will help our understanding of the quantum nature of BHs.

In practical terms, the construction of the HWF starts from the spectral decomposition of the quantum mechanical state that represents a matter source localized in space. By expressing the energy in terms of the gravitational (Schwarzschild) radius, as it would be classically determined according to the Einstein equations, the spectral decomposition then directly yields the HWF. The normalised HWF supplies the probability for an observer to detect a gravitational radius of a certain areal radius, centred around the source in the quantum state that was used in the first place. The gravitational radius can then be interpreted as a horizon if the probability of finding the particle inside it, is reasonably high. In other words, the horizon size is necessarily “fuzzy” in this QM description, just like the position of the particles that sources the geometry.

One would expect that quantum modifications of the sort mentioned above should be important only for small BHs, with masses close to the Planck scale, and that the quantum corrections to the classical solutions of large mass BHs should be negligible. A

large BH can have arbitrary small curvature near its horizon, and we have numerous tests of gravity in such low density regimes. However, BHs are very peculiar objects, in that they trap any signal inside the horizon *by definition*, no matter how weak tidal forces are in its neighbourhood, and it may happen that they are better described as macroscopic quantum states (see for instance, the Bose-Einstein Condensate model of refs. [10–14]).

Charged BHs were subject to many theoretical studies in the past [15]. In this work, we shall extend the HWF formalism to the Reissner-Nordström (RN) geometry and investigate the probability for an electrically charged source represented by a Gaussian wave-packet to be a BH, and to have an actual inner horizon. In the semiclassical approach, the latter is a Cauchy horizon and is sometimes associated with an instability known as “mass inflation”: any small matter perturbation will blue-shift unboundedly just outside this horizon, and inevitably produce a large deformation to the background geometry [16]. Although the existence of this effect is still debated (see, e.g. refs. [17–19]), it is clear from the classical causal structure of the RN geometry that, if there is matter falling through the outer horizon, it should accumulate outside the inner Cauchy horizon, and eventually lead to a large backreaction there. On the other hand, matter inside the inner horizon could escape the Cauchy horizon and produce a deformation as well. It is thus interesting to study under which conditions the inner horizon survives in the QM treatment.

2 The HWF formalism

The formalism introduced in refs. [5–7] is based on lifting the gravitational radius R_H of a QM system to the rank of a quantum operator. The coordinate r in a spherical metric is invariantly related to the geometrical area $4\pi r^2$ of the sphere centred on the origin $r = 0$, therefore being a suitable candidate for an observable in the quantum theory. Moreover, the horizon radius $r = R_H$ represents the location of trapping surfaces (surfaces where the escape velocity equals the speed of light) and thus determines the causal structure of the space-time, which again is likely an observable property in the quantum theory.

We remind the readers that in a neutral spherically symmetric system,

$$R_H(r) = 2\ell_p \frac{M(r)}{m_p}, \quad (2.1)$$

where ℓ_p is the Planck length and m_p the Planck mass,¹ and

$$M(r, t) = 4\pi \int_0^r \rho(\bar{r}, t) \bar{r}^2 d\bar{r} \quad (2.2)$$

is the Misner-Sharp mass. Here $M = M(r, t)$ represents the *total energy* inside the sphere of area $4\pi r^2$ (thus, roughly speaking, including the negative gravitational energy) and is related to the energy density ρ of the source via the *flat* space volume. A specific value of r is a trapping surface if $R_H(r) = r$, while, if $R_H(r) < r$, the gravitational radius is still well-defined but does not correspond to any causal surface.

¹We shall use units with $c = k_B = 1$, and always display the Newton constant $G_N = \ell_p/m_p$, so that $\hbar = \ell_p m_p$.

Going back to the HWF formalism, let us start from QM states representing *spherically symmetric* objects, which are *localized in space* and *at rest* in the chosen reference frame. Such particles are consequently described by wave-functions $\psi_S \in L^2(R^3)$, which we assume can be decomposed into energy eigenstates,

$$|\psi_S\rangle = \sum_E C(E) |\psi_E\rangle, \quad (2.3)$$

where the sum represents the spectral decomposition in Hamiltonian eigenmodes,

$$\hat{H} |\psi_E\rangle = E |\psi_E\rangle. \quad (2.4)$$

The actual Hamiltonian H needs not be specified yet.²

The expression of the Schwarzschild radius in eq. (2.1) can be inverted to obtain

$$M = m_p \frac{R_H}{2\ell_p}, \quad (2.5)$$

and remembering that M represents the total energy, it can be used to define the (unnormalized) “horizon wave-function” as

$$\tilde{\psi}_H(R_H) = C(m_p R_H/2\ell_p). \quad (2.6)$$

The normalisation of $\tilde{\psi}_H$ is finally fixed by means of the scalar product

$$\langle \psi_H | \phi_H \rangle = 4\pi \int_0^\infty \psi_H^*(R_H) \phi_H(R_H) R_H^2 dR_H, \quad (2.7)$$

again in agreement with the geometrical meaning of the variable R_H as yielding the area $4\pi R_H^2$ of the corresponding sphere. Our interpretation of the normalised ψ_H is then that it yields the probability for an observer to detect a horizon (necessarily “fuzzy”, like the position of the particle itself) of areal radius $r = R_H$, associated with the particle in the quantum state ψ_S .

In more details, starting from the wave-function ψ_H associated with ψ_S , we can now calculate the probability density for the particle to lie inside its own horizon of radius $r = R_H$:

$$\mathcal{P}_<(r < R_H) = P_S(r < R_H) \mathcal{P}_H(R_H), \quad (2.8)$$

where

$$P_S(r < R_H) = 4\pi \int_0^{R_H} |\psi_S(r)|^2 r^2 dr \quad (2.9)$$

is the probability that the particle is inside a sphere of radius $r = R_H$, and

$$\mathcal{P}_H(R_H) = 4\pi R_H^2 |\psi_H(R_H)|^2 \quad (2.10)$$

is the probability density that the sphere of radius $r = R_H$ is a horizon. Finally, the probability that the particle is a black hole will be obtained by integrating (3.13) over all possible values of the horizon radius, namely

$$P_{BH} = \int_0^\infty \mathcal{P}_<(r < R_H) dR_H. \quad (2.11)$$

²This is where, for instance, the self-gravity of the particle may enter.

To anticipate the calculations in the next section, we note that electrically charged BHs have two gravitational radii, and we will therefore define two corresponding operators, namely \hat{R}_\pm . We shall then define a HWF for each one, and obtain the probabilities for both the inner and outer horizon to exist.

3 Electrically charged spherical sources

We start from the RN metric, which can be written as

$$ds^2 = -f dt^2 + f^{-1} dr^2 + r^2 (d\theta^2 + \sin^2 \theta d\phi^2), \quad (3.1)$$

with

$$f = 1 - \frac{2\ell_p M}{m_p r} + \frac{Q^2}{r^2}, \quad (3.2)$$

where M and Q respectively represent the ADM mass and charge of the source. It is now convenient to introduce the specific charge

$$\alpha = \frac{|Q| m_p}{\ell_p M}. \quad (3.3)$$

The case $\alpha = 0$ reduces to the neutral Schwarzschild metric, and shall not be reconsidered here (see refs. [5–9]). For $0 < \alpha < 1$, the above metric contains two horizons, namely

$$\begin{aligned} R_\pm &= \ell_p \frac{M}{m_p} \pm \sqrt{\left(\ell_p \frac{M}{m_p}\right)^2 - Q^2} \\ &= \ell_p \frac{M}{m_p} \left(1 \pm \sqrt{1 - \alpha^2}\right), \end{aligned} \quad (3.4)$$

and represents a BH. The two horizons overlap for $\alpha = 1$, the so-called extremal BH case, while for $\alpha > 1$ no horizon exists and the central singularity is therefore accessible to an outer observer. This is the prototype of a naked singularity, but we shall not consider this case here (the corresponding HWF is the topic of ref. [20]).

We shall now investigate the case $0 < \alpha \leq 1$, which classically possesses at least one horizon, from a QM perspective. We first determine the HWFs and then calculate the probabilities for both the inner and outer horizon to exist. For this purpose, the classical relations (3.4) will be lifted to the rank of equations for the operators \hat{R}_\pm and \hat{M} , which are chosen to act multiplicatively on the HWF (with the specific charge α viewed as a simple parameter).

3.1 HWF for Gaussian source

The source for the RN space-time is taken to be an electrically charged massive particle at rest in the origin of the reference frame, represented by the spherically symmetric Gaussian wave-function

$$\psi_S(r) = \frac{e^{-\frac{r^2}{2\ell^2}}}{\ell^{3/2} \pi^{3/4}}. \quad (3.5)$$

The width of the Gaussian ℓ is assumed to be the minimum compatible with the Heisenberg uncertainty principle, that is

$$\ell = \lambda_m \simeq \ell_p \frac{m_p}{m}, \quad (3.6)$$

where λ_m is the Compton length of the particle of rest mass m (this assumptions is thoroughly investigated in ref. [7]). The spectral decomposition of eq. (3.5) is easily obtained from assuming the relativistic mass-shell relation in flat space,

$$M^2 = p^2 + m^2, \quad (3.7)$$

and by going to momentum space,

$$\psi_S(p) = \frac{e^{-\frac{p^2}{2\Delta^2}}}{\Delta^{3/2} \pi^{3/4}}, \quad (3.8)$$

where $p^2 = \vec{p} \cdot \vec{p}$ is the square modulus of the spatial momentum, and the width $\Delta = m_p \ell_p / \ell \simeq m$.

For $\alpha < 1$, one can now write a HWF for each of the two horizons in eq. (3.4). In fact, the total energy M can be expressed in terms of the horizon radii as

$$\ell_p \frac{\hat{M}}{m_p} = \frac{\hat{R}_+ + \hat{R}_-}{2}, \quad (3.9)$$

and

$$\hat{R}_{\pm} = \hat{R}_{\mp} \frac{1 \pm \sqrt{1 - \alpha^2}}{1 \mp \sqrt{1 - \alpha^2}}, \quad (3.10)$$

where M , R_+ , and R_- were promoted to operators \hat{M} , \hat{R}_+ , and \hat{R}_- related to the corresponding observables. It is important to remark that our choice is not unique, because of the usual ambiguities that emerge when going from a classical to the quantum formalism. The unnormalized HWFs for R_+ and R_- are obtained by expressing p from eq. (3.7) in terms of M in eq. (3.9), and then replacing one of the relations in eq. (3.10) into eq. (3.8). These manipulations yield

$$\begin{aligned} \psi_H(R_{\pm}) &= \mathcal{N}_{\pm} \Theta(R_{\pm} - R_{\min\pm}) \\ &\times \exp \left\{ -\frac{m_p^2 R_{\pm}^2}{2 \Delta^2 \ell_p^2 (1 \pm \sqrt{1 - \alpha^2})^2} \right\}. \end{aligned} \quad (3.11)$$

The step function in the first line above accounts for the minimum energy $M = m$ in the spectral decomposition of the wave-function (3.5), which corresponds to

$$R_{\min\pm} = \ell_p \frac{m}{m_p} \left(1 \pm \sqrt{1 - \alpha^2} \right). \quad (3.12)$$

Finally, the normalisations \mathcal{N}_{\pm} are fixed again by using the scalar product in eq. (2.7) for both R_{\pm} .³

³Explicit expressions of \mathcal{N}_{\pm} are very cumbersome and not particularly significant, thus shall be omitted throughout the paper.

The probability density that the particle lies inside its own horizons of size $r = R_{\pm}$ can now be calculated starting from the wave-functions (3.11) associated with (3.5) as

$$\mathcal{P}_{<\pm}(r < R_{\pm}) = P_S(r < R_{\pm}) \mathcal{P}_H(R_{\pm}), \quad (3.13)$$

where

$$P_S(r < R_{\pm}) = 4\pi \int_0^{R_{\pm}} |\psi_S(r)|^2 r^2 dr \quad (3.14)$$

is the probability that the particle is inside the sphere $r = R_{\pm}$, and

$$\mathcal{P}_H(R_{\pm}) = 4\pi R_{\pm}^2 |\psi_H(R_{\pm})|^2 \quad (3.15)$$

is the probability density that the sphere $r = R_{\pm}$ is a horizon. Finally, one can integrate (3.13) over all possible values of the horizon radius R_{\pm} to find the probability for the particle described by the wave-function (3.5) to be a BH, namely

$$P_{BH+} = \int_{R_{\min+}}^{\infty} \mathcal{P}_{<+}(r < R_{+}) dR_{+}. \quad (3.16)$$

The analogous quantity for R_{-} ,

$$P_{BH-} = \int_{R_{\min-}}^{\infty} \mathcal{P}_{<-}(r < R_{-}) dR_{-}, \quad (3.17)$$

will instead be viewed as the probability that the particle lies further inside its inner horizon. It is already clear from these definitions that $P_{BH-} < P_{BH+}$, and it is only when P_{BH-} is significantly close to one that we can say that both R_{-} and R_{+} are physically realised.

3.2 Inner and outer horizon probabilities

The probabilities $P_{BH\pm}$ can only be computed numerically, and we shall therefore display their behaviour graphically.

In the upper panel of figure 1 we show the probability density that the particle lies inside the outer horizon $r = R_{+}$ from eq. (3.13) for two values of the Gaussian width $\ell = \lambda_m \sim m^{-1}$ (above and below the Planck scale). This probability clearly decreases when m decreases below the Planck mass, which corresponds to ℓ increasing above ℓ_p . A similar analysis is presented in the lower panel of the same figure for the probability density that the particle lies inside the inner horizon $r = R_{-}$. It is obvious that the probabilities in the second case are much smaller with decreasing α . As expected, the two probability densities are identical for the extremal case $\alpha = 1$ (thick and thin dashed lines), since the two horizons coincide.

The probabilities $P_{BH\pm}$ are obtained by performing the integrations (3.16) and (3.17). The plot in figure 2 shows these probabilities as functions of α for values of the particle mass above, equal to and below the Planck mass. When analyzing the outer horizon, one notices that P_{BH+} stays very close to one for mass values larger than the Planck scale. However, for $m \lesssim m_p$ (when the width of the Gaussian wave-packet $\ell \gtrsim \ell_p$), P_{BH+} clearly decreases as the BH specific charge increases to one. Note that this probability is

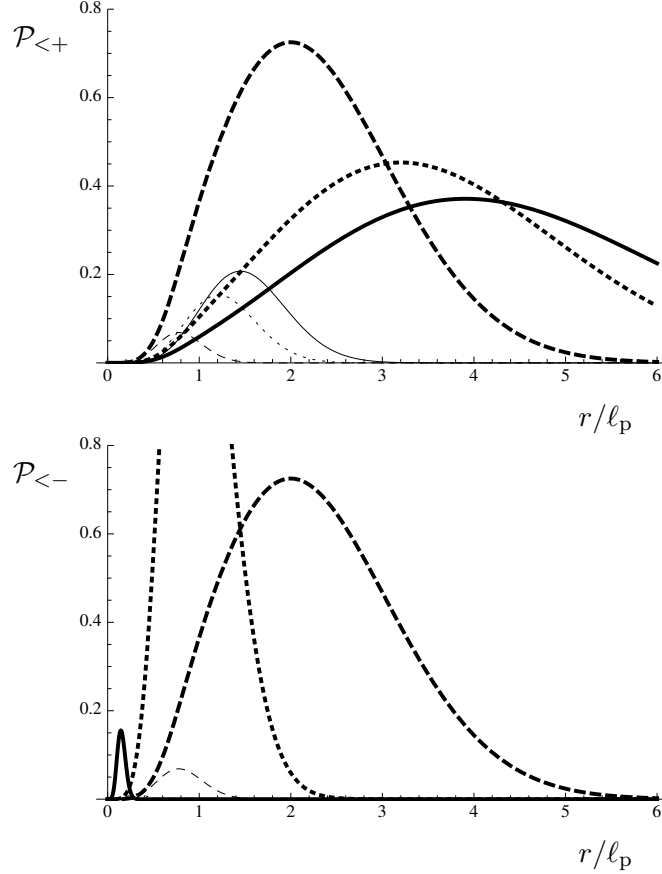


Figure 1. Top panel: probability density $\mathcal{P}_{<+}$ in eq. (3.13) that the particle is inside its outer horizon $r = R_+$, for $m = 2m_p$ (thick lines) and $m = 0.5m_p$ (thin lines) with $\alpha = 0.3$ (continuous lines), $\alpha = 0.8$ (dotted lines) and $\alpha = 1$ (dashed lines). Bottom panel: probability density $\mathcal{P}_{<-}$ in eq. (3.13) that particle is inside its inner horizon $r = R_-$, for $m = 2m_p$ (thick lines) and $m = 0.5m_p$ (thin lines) with $\alpha = 0.3$ (continuous lines), $\alpha = 0.8$ (dotted lines) and $\alpha = 1$ (dashed lines). For $\alpha = 1$, the two horizons coincide and $\mathcal{P}_{<-} = \mathcal{P}_{<+}$.

not exactly zero even for values of the mass smaller than m_p . For instance, in the case $m = 0.5m_p$, corresponding to a width $\ell = 2\ell_p$ of the Gaussian wave-packet, $P_{\text{BH}+} \simeq 0.2$ for a considerable range of values of α . It only decreases below 0.1 when α approaches one, therefore when the BH becomes maximally charged. The situation is very different for the inner horizon. The same plot shows that the probability $P_{\text{BH}-}$ starts from almost zero for small values of the charge-to-mass ratio and increases with α . The larger the mass of the particle, the smaller the value of α for which the probability starts to become significant. Still, there is a considerable range of values of the specific charge for which, while $P_{\text{BH}+} \simeq 1$ thus making the object a BH, the probability for the inner horizon to exist is approximately zero.

Figure 3 shows the probabilities $P_{\text{BH}\pm}$ as functions of the mass m for $\alpha = 0.3, 0.8$ and 1 . From this plot it becomes clear that for smaller values of α , the probability $P_{\text{BH}+}$ starts to increase from zero to one at smaller values of m . The opposite is true when analyzing $P_{\text{BH}-}$.

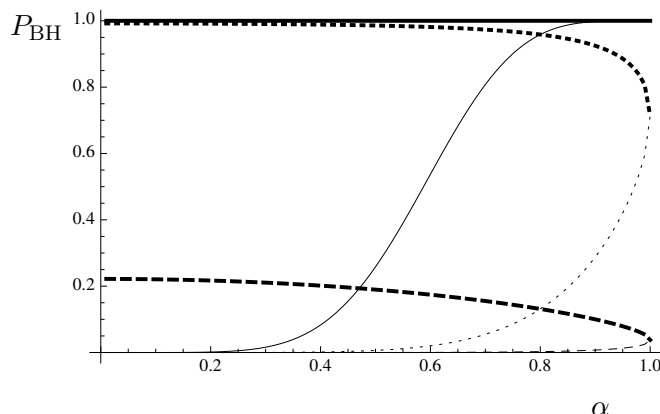


Figure 2. Probability $P_{\text{BH}+}$ in eq. (3.16) for the particle to be a BH (thick lines) and $P_{\text{BH}-}$ in eq. (3.17) for the particle to be inside its inner horizon (thin lines) as functions of α for $m = 2m_p$ (continuous line), $m = m_p$ (dotted line) and $m = 0.5m_p$ (dashed line). For $\alpha = 1$ the two probabilities merge.

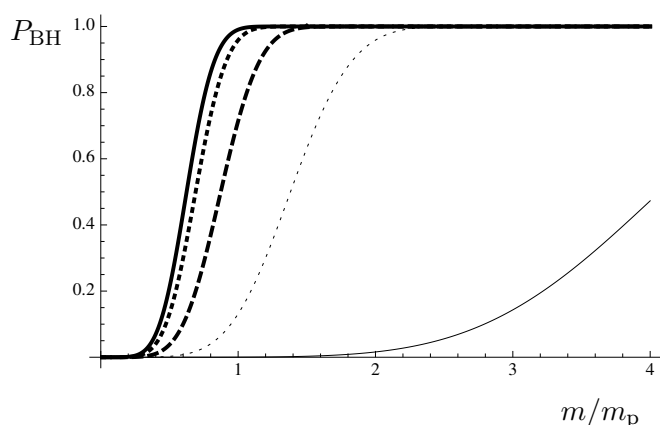


Figure 3. Probability $P_{\text{BH}+}$ in eq. (3.16) for the particle to be a BH (thick lines) and $P_{\text{BH}-}$ in eq. (3.17) for the particle to be inside its inner horizon (thin lines) as functions of the mass for $\alpha = 0.3$ (continuous line), $\alpha = 0.8$ (dotted line) and $\alpha = 1$ (dashed line). For $\alpha = 1$ thick and thin dashed lines overlap.

For the smallest value of the charge-to-mass ratio considered here, $\alpha = 0.3$, it is only around a particle mass of $m \simeq 6m_p$ that both probabilities $P_{\text{BH}+}$ and $P_{\text{BH}-}$ have values close to one, while $P_{\text{BH}+}$ already increases to one around m_p . This means that in the range of masses between m_p and $6m_p$, the probability $P_{\text{BH}+} \simeq 1$ while $P_{\text{BH}-} \simeq 0$. This mass range increases even more for smaller values of the specific charge, but it decreases to zero in the maximally charged limit.

Our main finding is therefore that there exists a considerable parameter space for m (around the Planck scale) and $\alpha < 1$ in which

$$P_{\text{BH}+} \simeq 1 \quad \text{and} \quad P_{\text{BH}-} \simeq 0. \quad (3.18)$$

Whether the particle is a BH or not is dictated by the existence of the outer horizon, therefore we interpret $P_{\text{BH}+} \simeq 1$ as meaning that the particle is (most likely) a BH.

However, the presence of an inner horizon at $r = R_-$ is important in light of the “mass inflation” instability and peculiar features of Cauchy horizons. Eq. (3.18) therefore means that the particle is (most likely) a BH, but no such peculiarities are expected to occur.

3.3 Generalised uncertainty principle

The uncertainty in the horizon size for a neutral BH was already investigated previously. In particular, it was shown to grow linearly with the BH mass and lead to a generalised uncertainty principle (GUP) in refs. [5, 6]. We can here repeat the same arguments for the outer horizon of the charged BH, and obtain similar results.

We first note that the expectation value

$$\begin{aligned}\langle \hat{R}_+ \rangle &= 4\pi \int_{R_{\min+}}^{\infty} |\psi_H(R_+)|^2 R_+^3 dR_+ \\ &= \frac{4 \left(1 + \sqrt{1 - \alpha^2}\right)}{2 + e\sqrt{\pi} \operatorname{erfc}(1)} \frac{\ell_P^2}{\ell} \\ &= R_+(\bar{M}),\end{aligned}\tag{3.19}$$

reproduces exactly the classical expression of R_+ in eq. (3.4) for $\ell = \lambda_m \sim m^{-1}$ and $\bar{M} = 4m/[2 + e\sqrt{\pi} \operatorname{erfc}(1)] \simeq 1.45m$.⁴ From

$$\begin{aligned}\langle \hat{R}_+^2 \rangle &= 4\pi \int_{R_{\min+}}^{\infty} |\psi_H(R_+)|^2 R_+^4 dR_+ \\ &= \frac{\left(1 + \sqrt{1 - \alpha^2}\right)^2 (10 + 3e\sqrt{\pi} \operatorname{erfc}(1))}{2(2 + e\sqrt{\pi} \operatorname{erfc}(1))} \frac{\ell_P^4}{\ell^2} \\ &\simeq R_+^2(\bar{M}),\end{aligned}\tag{3.20}$$

one can then calculate the uncertainty

$$\Delta R_+ = \sqrt{\langle \hat{R}_+^2 \rangle - \langle \hat{R}_+ \rangle^2} \simeq R_+ \sim m,\tag{3.21}$$

which, like in the neutral Schwarzschild case, grows linearly with the mass m of the source. This signals the fact that the state of such objects would remain QM even in an astrophysical regime, where we instead expect the horizon has a sharp location, and supports alternative models of large BHs, such as the ones in refs. [10–14].

If we now combine the horizon uncertainty (3.21) with the usual QM uncertainty in the radial size of the source,

$$\begin{aligned}\Delta r^2 &= 4\pi \int_0^{\infty} |\psi_S(r)|^2 r^4 dr - \left(4\pi \int_0^{\infty} |\psi_S(r)|^2 r^3 dr\right)^2 \\ &\simeq \ell^2,\end{aligned}\tag{3.22}$$

⁴This mass renormalisation, with $\bar{M} > m$, can be easily understood by noting that the source wavefunction ψ_S contains energy contributions from momenta $p > 0$.

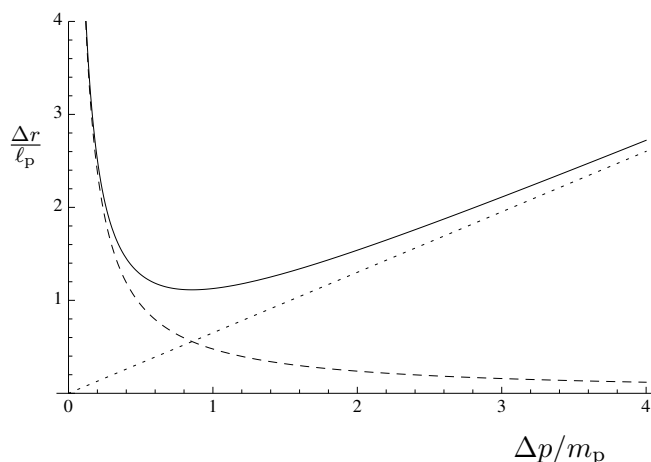


Figure 4. Uncertainty relation (3.23) (solid line) as a combination of the QM uncertainty (dashed line) and the uncertainty in horizon radius (dotted line), for $\gamma = 1$.

we can finally obtain a total uncertainty

$$\begin{aligned}\Delta r &\equiv \sqrt{\langle \Delta r^2 \rangle} + \gamma \sqrt{\langle \Delta R_+^2 \rangle} \\ &\simeq \ell_p \frac{m_p}{\Delta p} + \gamma \ell_p \frac{\Delta p}{m_p},\end{aligned}\tag{3.23}$$

where γ is a coefficient of order one. The result is plotted in figure 4, where it is also compared to the usual Heisenberg uncertainty in the size Δr of a state with an uncertainty in momentum given by

$$\begin{aligned}\Delta p^2 &= 4\pi \int_0^\infty |\psi_S(p)|^2 p^4 dp - \left(4\pi \int_0^\infty |\psi_S(p)|^2 p^3 dp\right)^2 \\ &\simeq \Delta^2 \simeq \ell^{-2}.\end{aligned}\tag{3.24}$$

We can therefore conclude that the outer horizon behaves qualitatively like the neutral Schwarzschild radius.

4 Conclusions and outlook

In this work we transparently applied some basic QM methods to the case of the RN BH. In our formalism, the location of the horizon is not given by a sharp classical value, instead it is described by a quantum wave-function with associated uncertainties. In addition, one can define a quantity which corresponds to the probability for a horizon to be formed. Since the RN BH has the inner and outer horizon, the structure of a quantum BH might be quite different from the structure of the classical one.

We first note here that, if one models a BH as a *point-like* source (i.e. one very narrow Gaussian wave-function), the uncertainty in the location of the outer horizon ΔR_+ grows with the mass of the source (since $\ell \sim m^{-1}$). This is in agreement with many postulated GUPs in the presence of gravity [21, 22], but implies that when the source is of astrophysical

size, the fluctuations become unacceptably large. This might imply that there must be a regime where the GUP gets replaced by the standard (non-gravitational) uncertainty principle of QM. In other words, the parameter γ in eq. (3.23) becomes very small in the semi-classical regime.

The other (somewhat related) option is that the single Gaussian source (3.5) does not appear to be a sensible model for large BHs. The fact that the uncertainty in the horizon location does not decrease for increasing rest mass m , might imply that a semiclassical behavior cannot be recovered at all in such a setup [5, 6]. This result was remarked in ref. [9], where it was then shown that modelling a BH as a large number of light constituents [10–14] does not suffer of this limitation. Because of that, we cannot here extend our findings to arbitrarily large BH mass straightforwardly.

We also note that modelling the source with some shape different from the Gaussian will not make qualitative differences. For example, using a step function will give a result in qualitative agreement with the Gaussian of similar width. In fact, one can reliably approximate any localised state ψ_S with a superposition of Gaussians. Given the linearity of the formalism, a superposition of, say, N Gaussians of roughly similar mass m will lead to a total HWF given by a superposition of the corresponding HWFs. In particular, if the width ℓ of the Gaussians are very narrow ($\ell \leq \ell_p$), we can understand what happens by simply replacing the superposition with one Gaussian of same ℓ and mass equal to Nm . The latter will increase $\langle \hat{R}_H \rangle \gg \ell$, and also increase the probability P_{BH} . However, if $\ell \sim 1/Nm$, the uncertainty $\Delta R_H \sim \langle \hat{R}_H \rangle$ and the system will never look classical. In contrast, if we model the source as N Gaussians of large width $\ell \gg \ell_p$, but $m \ll m_p$, like in the Bose-Einstein Condensate model of Dvali and Gomez [10–14], then $\langle \hat{R}_H \rangle \sim R_H$ and $\Delta R_H \ll \langle \hat{R}_H \rangle$ for large N , which does instead look classical. If the source has electrical charge, the above argument then holds for R_+ . Since the N constituents of the condensate have width $\ell \sim R_+$, it immediately follows that the inner horizon R_- will have very small probability to exist (expect perhaps in a near-extremal configuration).

To conclude, we can speculate that, at least in a quantum regime of BH masses, say $m \lesssim 10m_p$, quantum fluctuations around the inner horizon are strong enough to prevent the instability expected according to the semiclassical analysis. More generally, the probability that any instability occurs will be as small as the probability $P_{\text{BH-}}$ that the source is located inside the inner horizon.

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